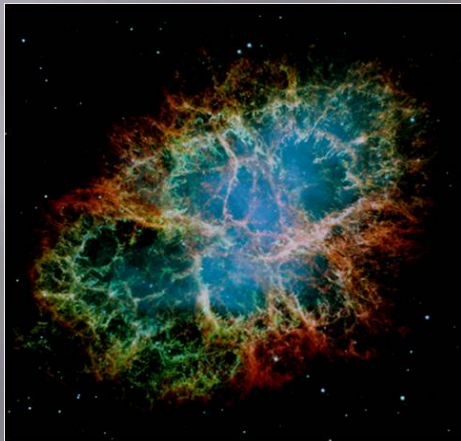
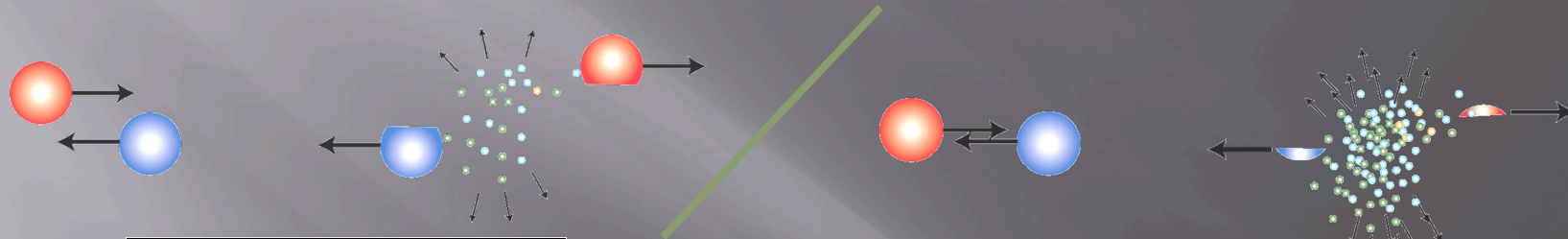


HOW TO PROBE EQUATION OF STATE OF NEUTRON-RICH NUCLEAR MATTER



Department of Physics
Kyoto University
Tetsuya MURAKAMI



Nuclear Equation of State

$$E(\rho, \delta) = E(\rho, 0) + E_{\text{sym}}(\rho)\delta^2 + o(\delta^4)$$

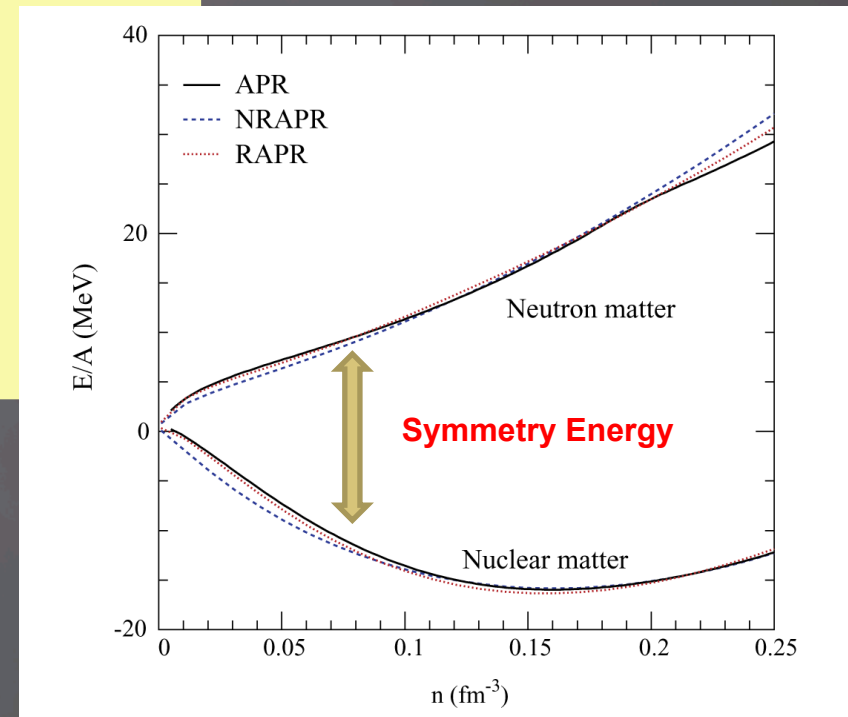
$$\delta = (\rho_n - \rho_p) / \rho$$

$$E_{\text{sym}}(\rho) = E_{\text{sym}}(\rho_0) + L\varepsilon + \frac{K_{\text{sym}}}{2}\varepsilon^2 + o(\varepsilon^3)$$

$$\varepsilon = (\rho - \rho_0) / 3\rho_0$$

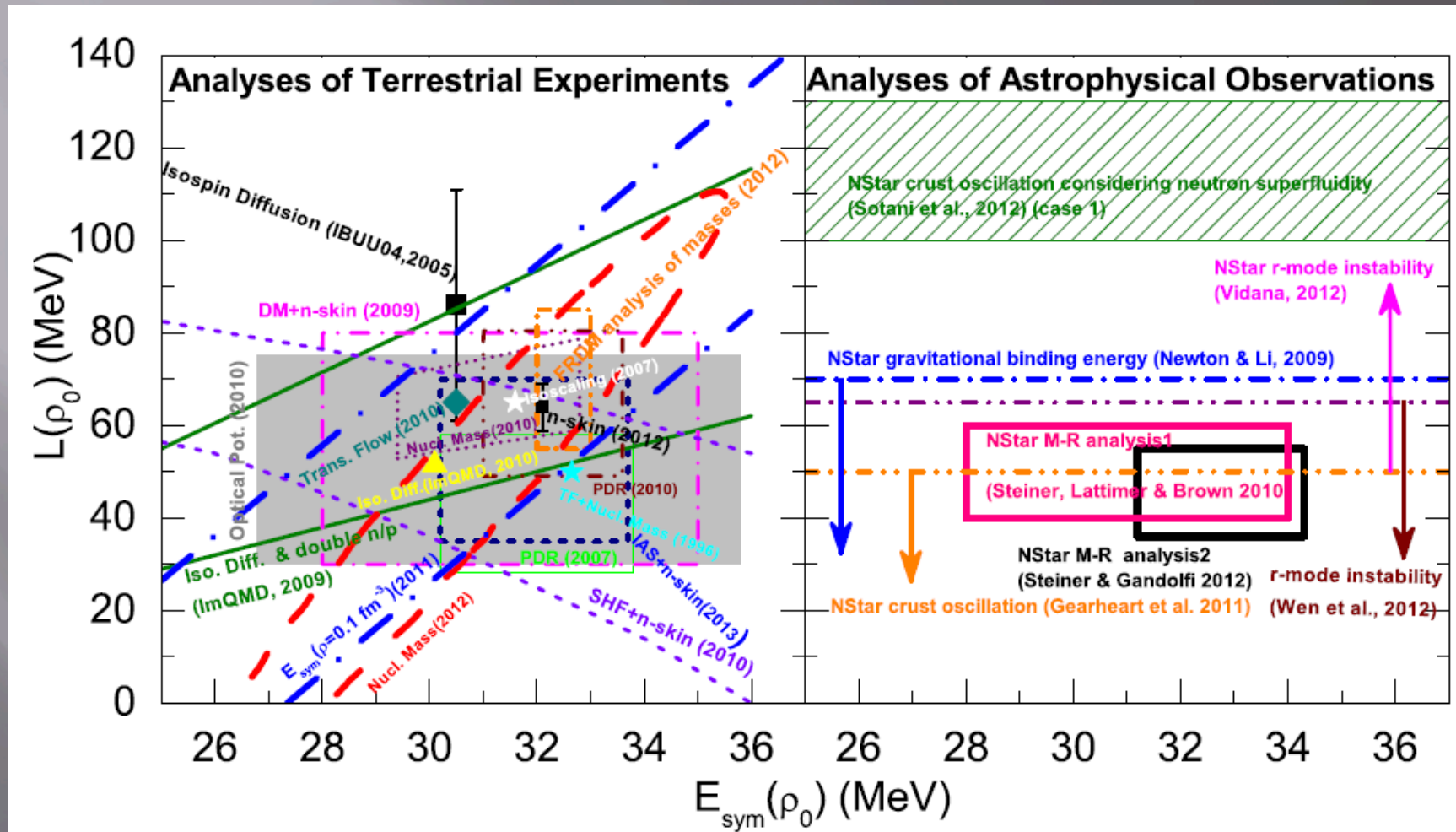
$$S_0 = E_{\text{sym}}(\rho_0) \Rightarrow J; \text{ sometimes}$$

$$L = 3\rho_0 \left. \frac{\partial E_{\text{sym}}(\rho)}{\partial \rho} \right|_{\rho=\rho_0} = (3/\rho_0)P_0$$



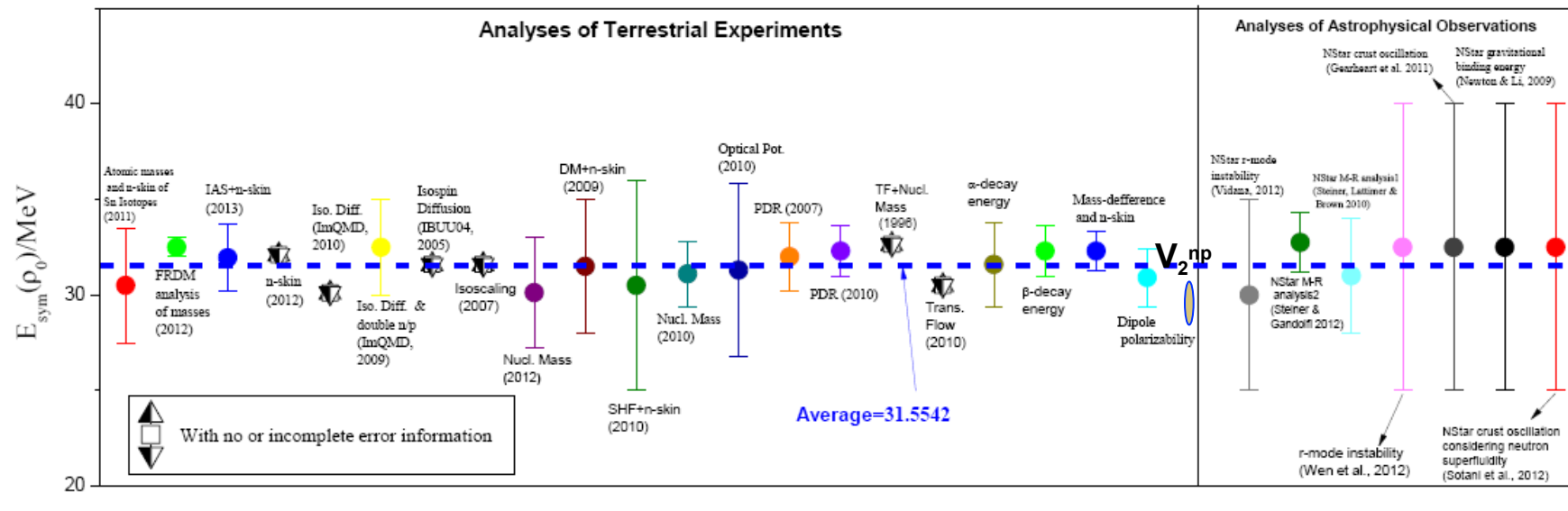
A.W. Steiner et al., Phys. Reports 411 (2005) 325

From summary talk at NuSYM13 by Bao-An Li (next two slides)



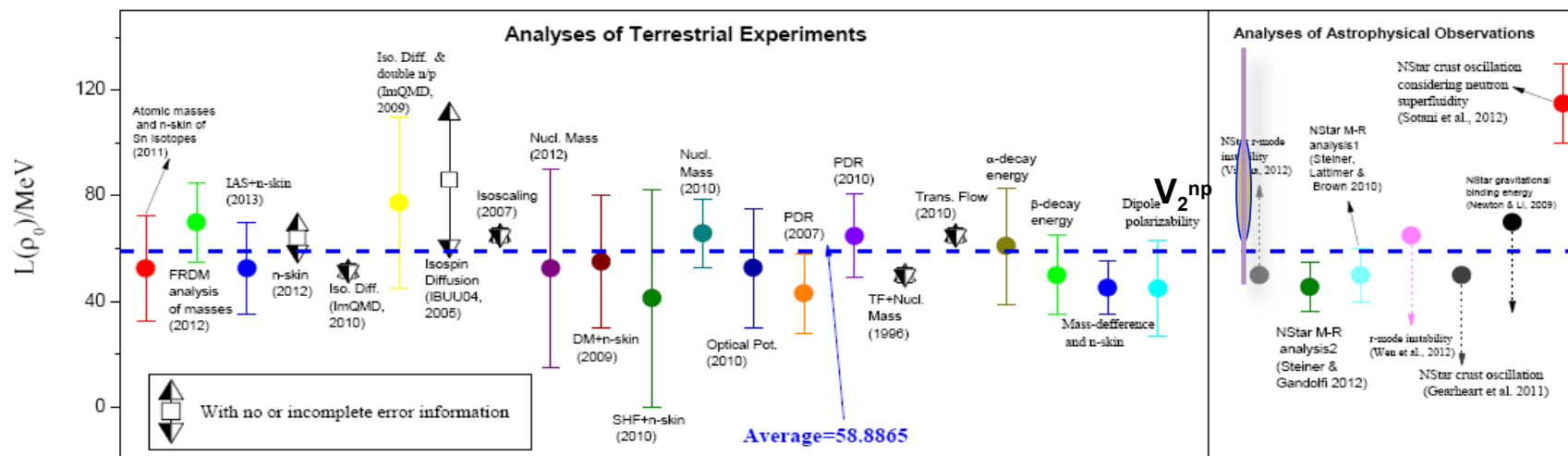
Workshop on Equation of State for Neutron
Star Matter

Nusym13 constraints on $E_{\text{sym}}(\rho_0)$ and L based on 29 analyses of some data



	E_{sym}	L
average of the means	31.55415	58.88646
standard deviation	0.915867	16.52645

Currently impossible to estimate a physically meaningful error bar



Quick Review of Each Approach Related “Our Innovative Area”

2013/10/25

Workshop on Equation of State for
Neutron Star Matter

FRDM: mass formula

P. Moller et al. ; Phys. Rev. Lett 108 (2012) 052501

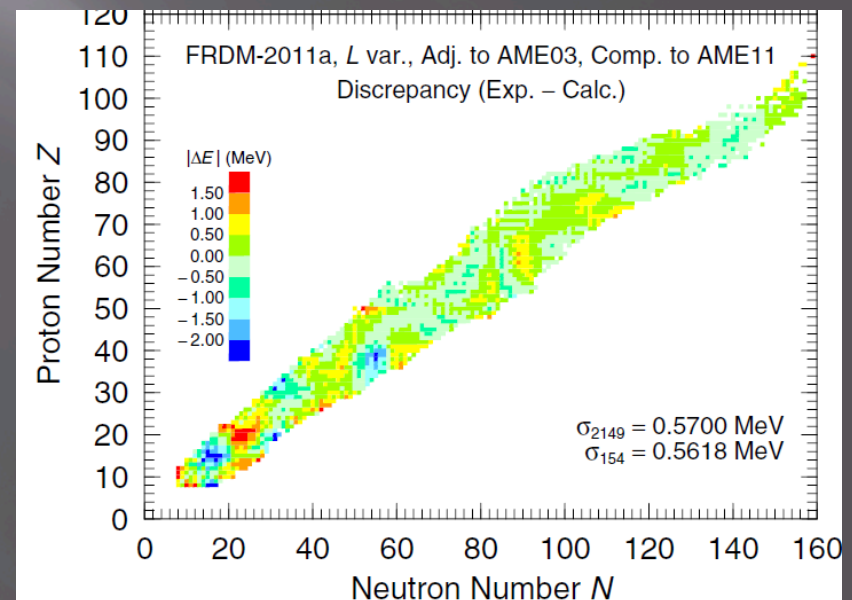
Conventional liquid drop model

$$BE = -a_1 A + a_2 A^{2/3} + a_3 (Z^2/A^{1/3}) + a_4 (N-Z)^2/A + \delta E$$

New approach using **Finite Range Droplet Model**
incorporated with the density dependence of symmetry
energy can provides accurate binding energies better
than 0.2% for all mass available in **AME10**.

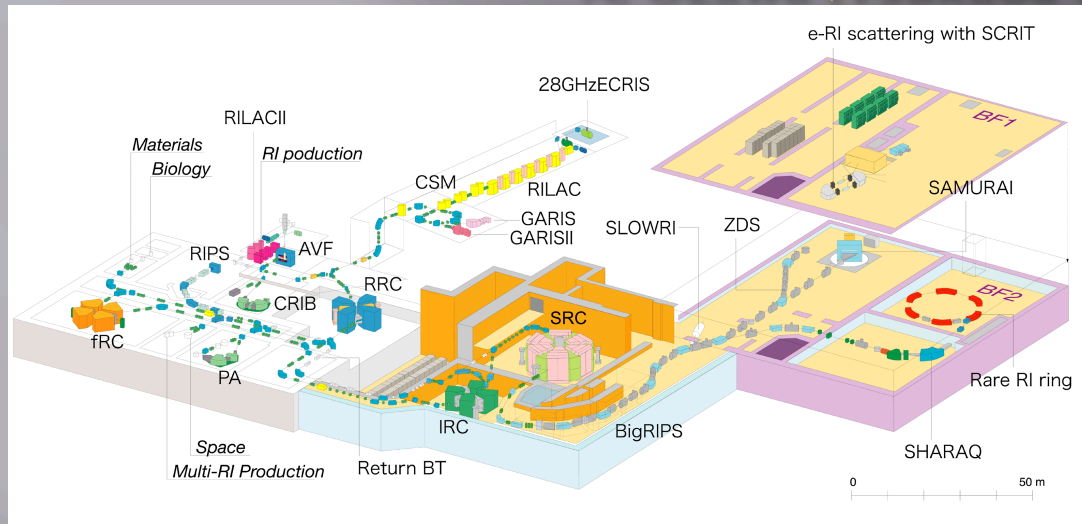


$$S_0 = 32.5 \pm 0.5 \text{ MeV and} \\ L = 70 \pm 15 \text{ MeV}$$

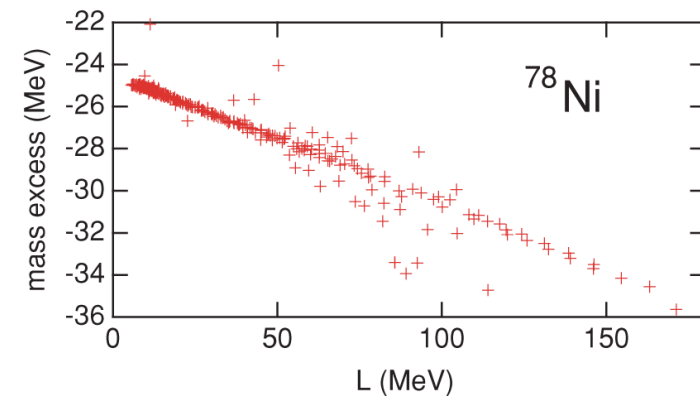


Future Direction of Mass Measurement:

A. Ozawa (Tsukuba Univ.) Group



Measure mass of ^{78}Ni with an accuracy of 10^{-6} to determine L

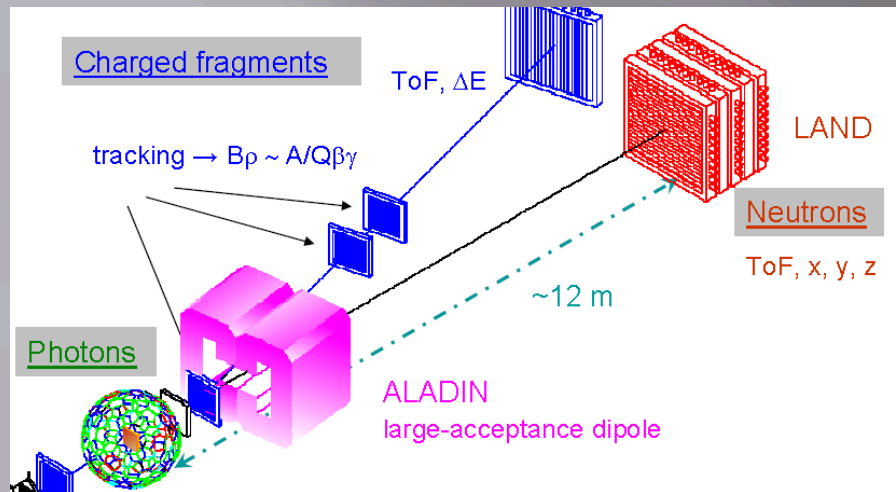


K. Oyamatsu and K. Iida, Phys. Rev. C 81 (2010) 054302

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Star Matter

Pygmy Dipole Resonances



From ^{68}Ni and ^{132}Sn data

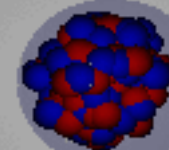
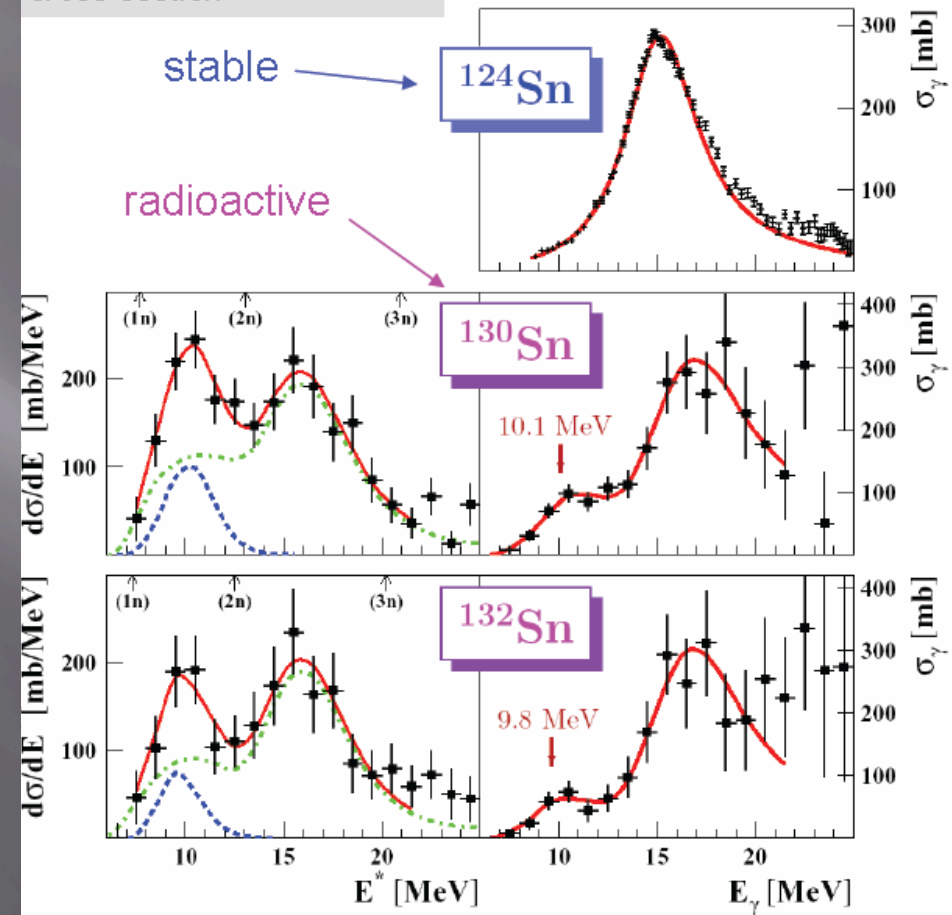
$L = 64.8 \pm 15.7 \text{ MeV}$ and

$S_0 = 32.3 \pm 1.3 \text{ MeV}$

(Carbone et al. PRC 81, 041301 (2010))

Electromagnetic-excitation
cross section

Photo-neutron cross section



SAMURAI Magnet



Superconducting **A**naluzer for **M**ulti particles from
Radio **I**sotope Beams

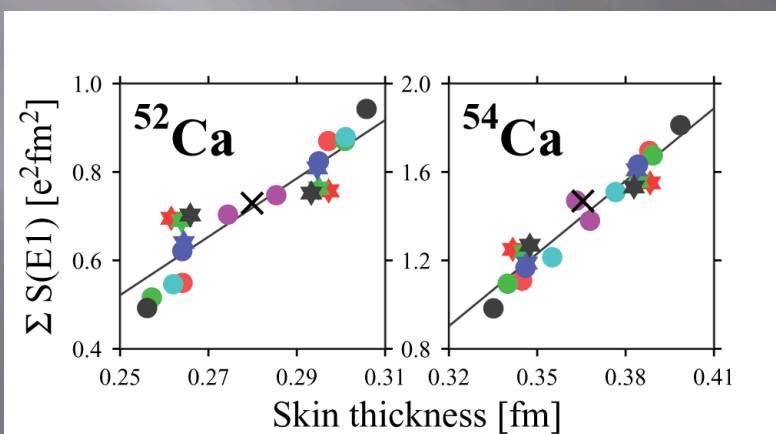
Future Direction:

CATANA for SAMURAI (T. Togano (TIT) et al.)

Isovector dipole response of $^{48-52}\text{Ca}$

→ skin thickness, EOS

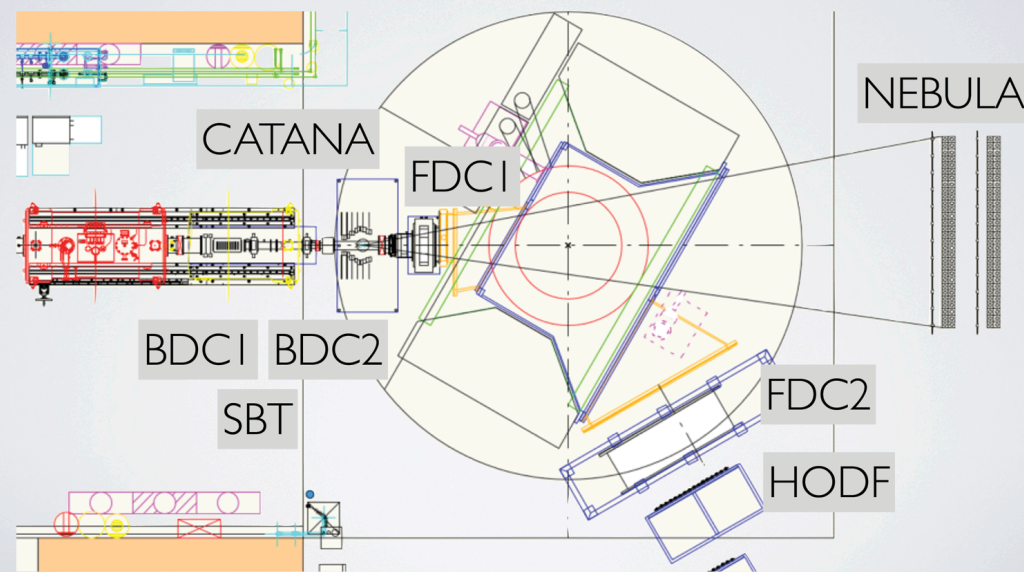
PDR strength and
dipole polarizability



T. Inakura et al., arXiv:1306.3089

Experimental setup

- Target: 500 mg/cm² Pb, 300 mg/cm² C



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Star Matter

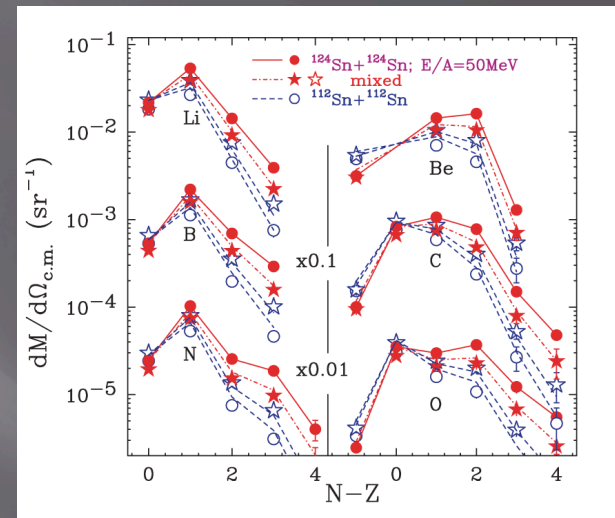
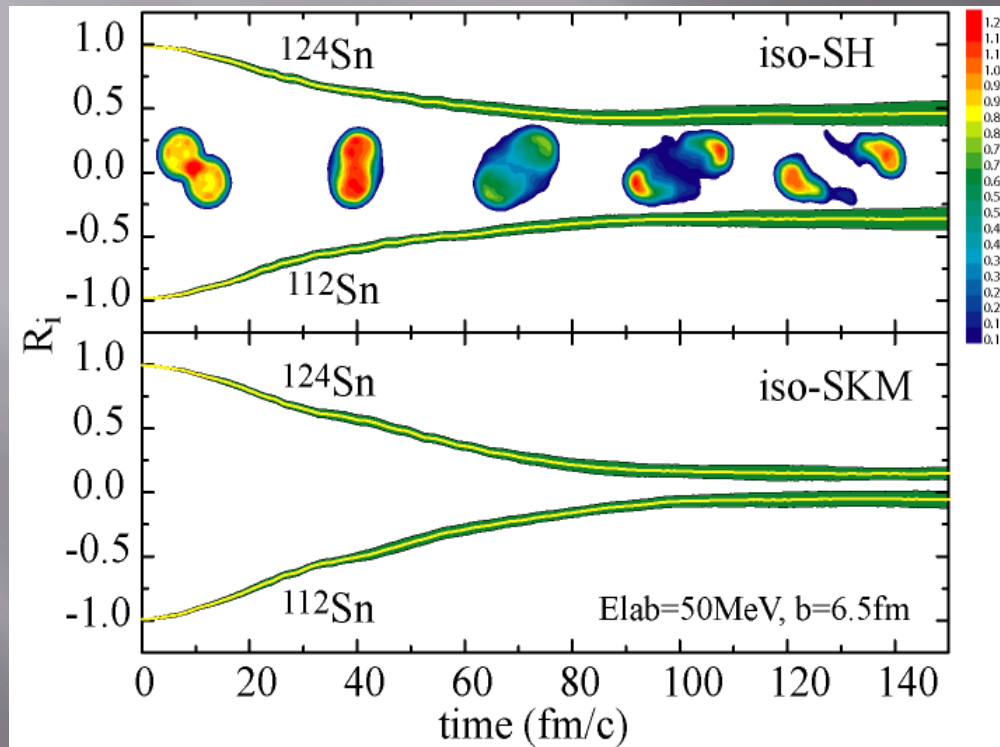
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Heavy Ion Collisions

compared to transport theory calculation
NSCL/MSU group

□ Isospin diffusion in peripheral collisions

Tsang et al., PRL 92 (2004) 062701



$$R_{21}(N, Z) = \frac{Y_2(N, Z)}{Y_1(N, Z)} \approx C \exp(N\alpha + Z\beta)$$

$$R_i(X) = 2 \frac{X - (X_{A+A} + X_{B+B})/2}{X_{A+A} - X_{B+B}}$$

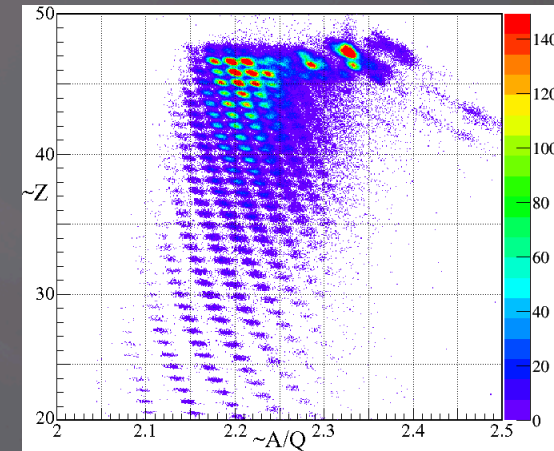
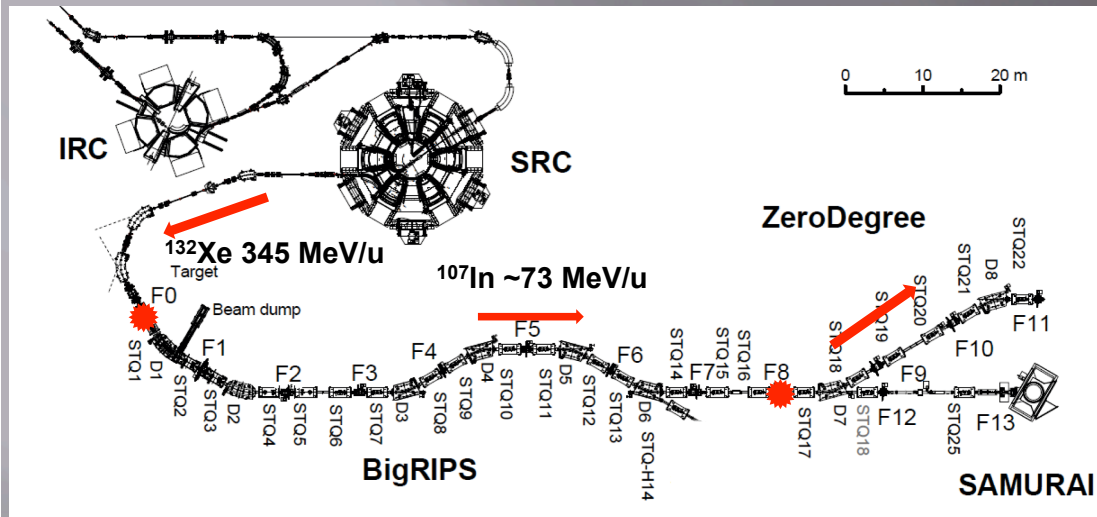
□ Yield ratio of neutron and proton

Zhang et.al., Phys. Lett. B 664 (2008) 145

Workshop on Equation of State for Neutron Star
Matter

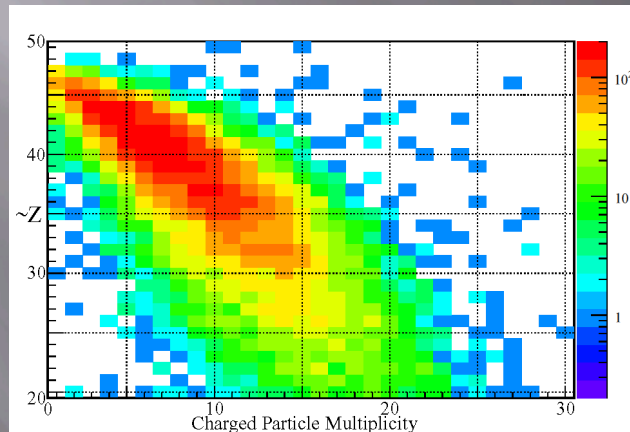
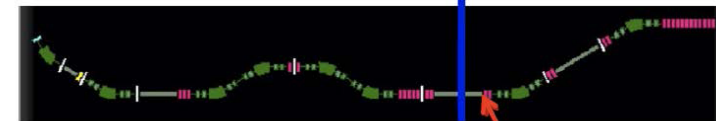
Extention at RIBF

MSU/TAMU/RIKEN/Kyoto Group in June 2013



BigRIPS

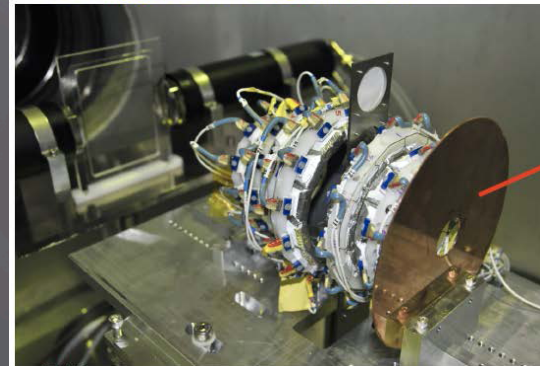
Zero Degree Spectrometer



$^{107}\text{In}(\text{beam}, \sim 73\text{MeV/u}) + ^{124}\text{Sn}(\text{Target})$

$^{112}\text{Sn}(\text{beam}, \sim 73\text{MeV/u}) + ^{112}\text{Sn}(\text{Target})$

Washington University microball

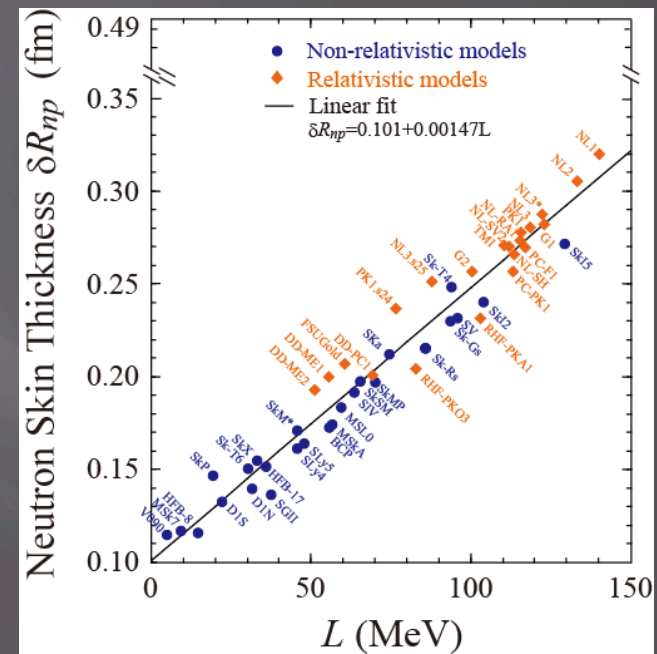


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Star Matter

From neutron skin thickness for Pb

- ▣ Parity violation electron scattering
Pb Radius Experiment (PREx) at Jefferson Lab.
 $\Delta R_{np} = 0.33^{+0.16}_{-0.18}$ fm
(S. Abrahamyan et al. PRL 108, 112502 (2012))
- ▣ Dipole Polarizability and PDR
 $\Delta R_{np} = 0.156^{+0.025}_{-0.021}$ fm (Tamii et al. PRL 107, 062502 (2011))
 $= 0.194 \pm 0.024$ fm (Carbone et al. PRC 81, 041301 (2010))
- ▣ Antiproton Atoms
 $\Delta R_{np} = 0.15 \pm 0.02$ fm (Trzcinska et al. PRL 87, 082501 (2001))
- ▣ Proton elastic scattering
 $\Delta R_{np} = 0.211^{+0.054}_{-0.063}$ fm (Zenihiro et al. PRC82, 054607 (2010))



X. Roca-Maza et al., *PRL*106, 252501 (2011)

Recoil Proton Spectrometer (RPS)

Transfer region	
Transfer region	$p = 20 - 120 \text{ MeV}, \Delta\Omega \sim 10 \text{ msr/deg.}$
Drift chamber	$436 \times 436 \text{ mm}^2 \text{ (x-y-x'-y')}$
Drift chamber	$440 \times 440 \text{ mm}^2 \times 2 \text{ mm}^t$
Calorimeter	$450 \times 50 \text{ mm}^2 \times 50 \text{ mm}^t$
Hydrogen	$\Phi 30 - 1 \text{ mm}^t$

transfer region

$p = 20 - 120 \text{ MeV}, \Delta\Omega \sim 10 \text{ msr/deg.}$

Drift chamber	$436 \times 436 \text{ mm}^2 \text{ (x-y-x'-y')}$
Drift chamber	$440 \times 440 \text{ mm}^2 \times 2 \text{ mm}^t$
Calorimeter	$450 \times 50 \text{ mm}^2 \times 50 \text{ mm}^t$
Hydrogen	$\Phi 30 - 1 \text{ mm}^t$

Plastic scintillator

Beam Drift Chamber(BDC1)

BDC2

BDC3

Solid hydrogen

veto scintillator

Plastic scintillator

CH₂ - C

Plastic scintillator

Si-Telescope

$L \sim 1 \text{ m}$

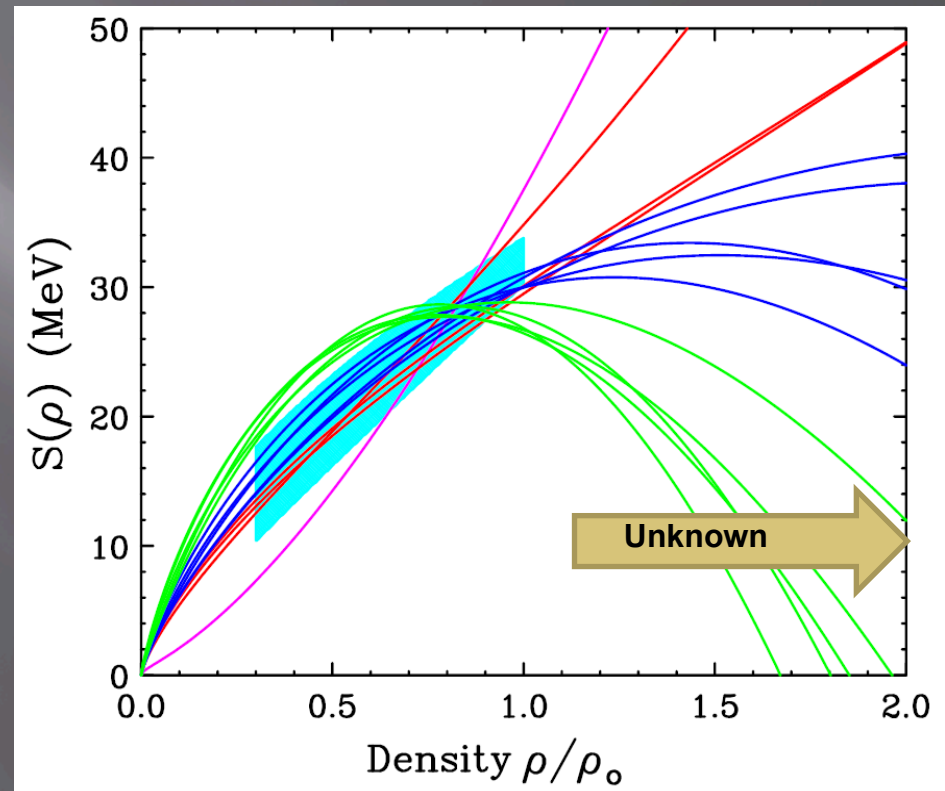
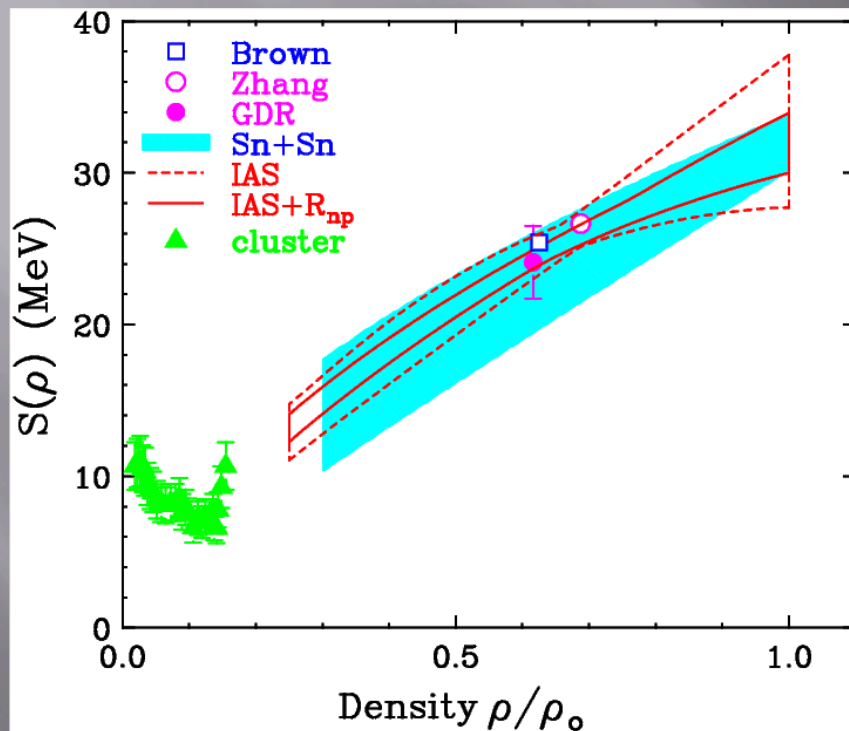
$L \sim 0.25 \text{ m}$

- ## Si-Telescope



2013/10/25

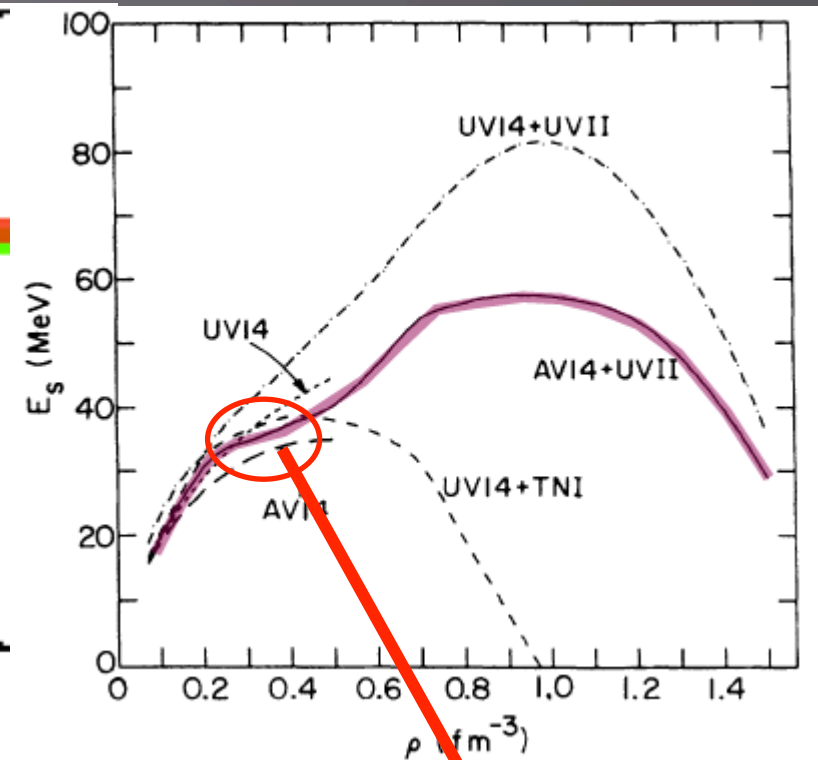
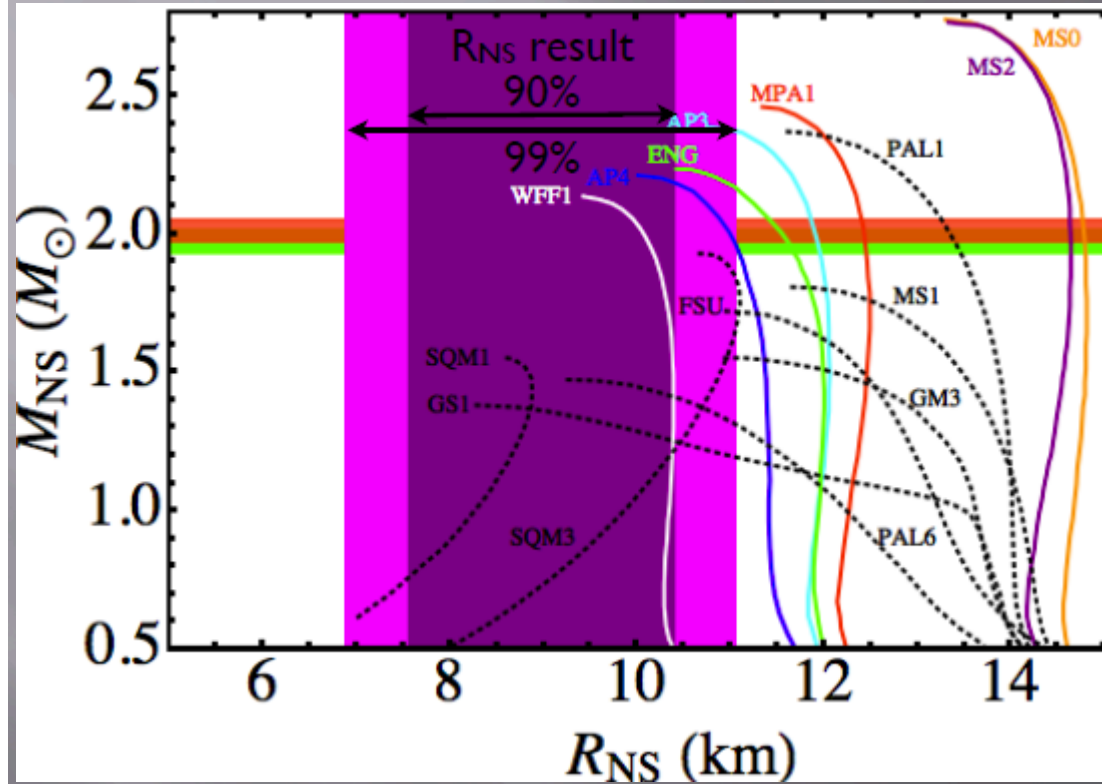
Updated Constraints from NuSYM13



Surprising Astrophysical Observation!!

Neutron star (Rutledge, Gulliot et al.)
APJ 2013 to be published

AV14+UVII ;
Wiringa, Fiks, & Fabrocini ,
Phys. Rev. C 38, 1010 (1988)



Observation:

$$M_{NS} \sim 2M_{\text{sun}}$$

$$R_{NS} \sim 9 \text{ km}$$

Equation of State
softening EoS at $\rho \sim 2\rho_0$
stiff EoS at high ρ

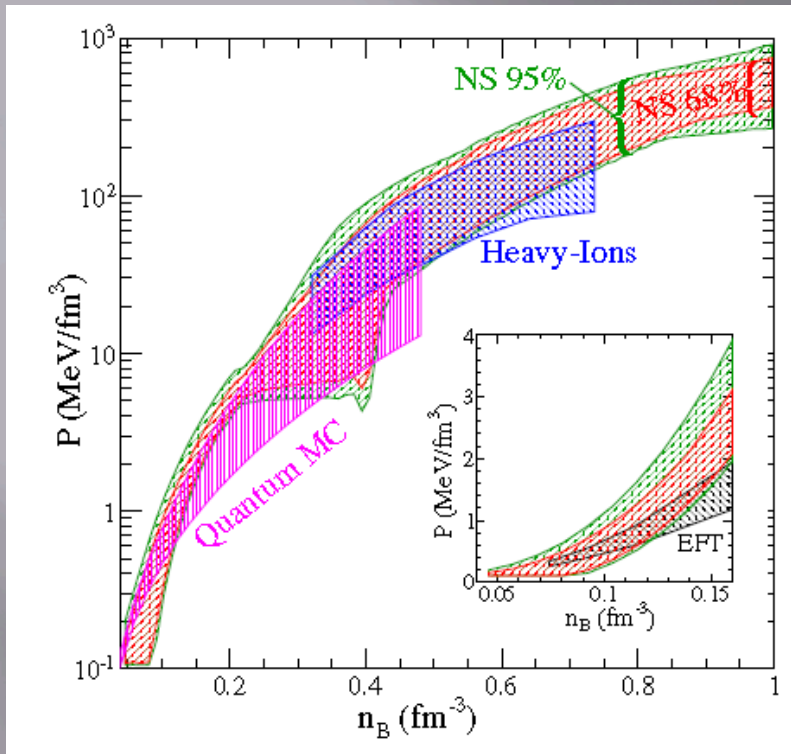
Workshop on Equation of State for Neutron Star
Matter

How to approach high density?

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Workshop on Equation of State for
Neutron Star Matter

Motivation of SPiRIT



A.W.Steiner et al. APJ 765, L5 (2013)

- ▣ Constraints from nuclear structure, nuclear reactions and neutron stars .
- ▣ Neutron star constraints at high density obtained from X-ray burst light curves.
- ▣ QMC constraints extrapolate nuclear structure information to high density. Width is compatible with uncertainty expected for a successful PREX experiment.
- ▣ Heavy ion collision constraints combine those from collective flow with guesses about symmetry energy. >> Still uncertain!!

Goal is to decrease the factor of 2 uncertainty in symmetry pressure at $\rho \approx 2\rho_0$.

Prediction of Bao-An Li

NPA 708 (2002) 365

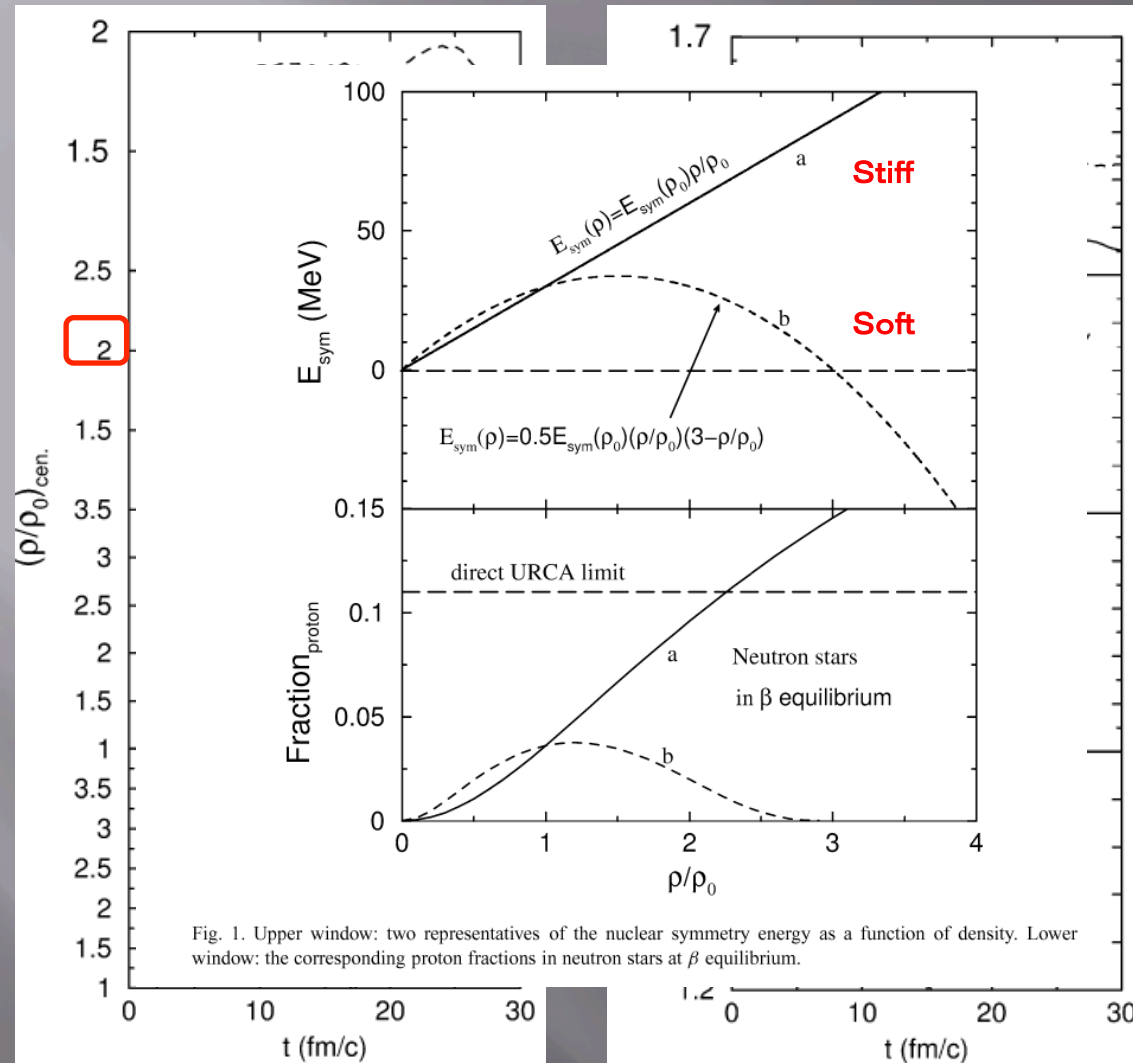
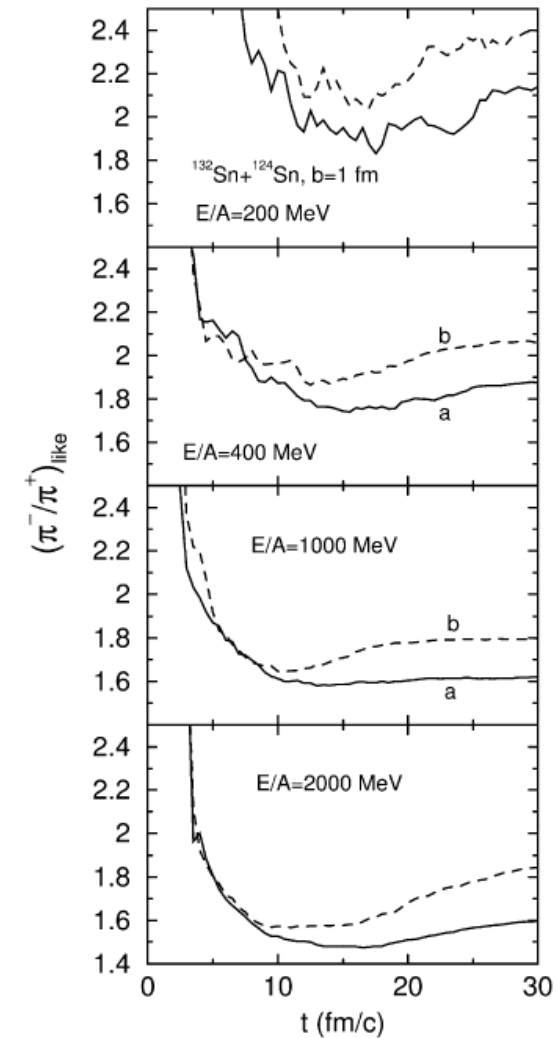


Fig. 1. Upper window: two representatives of the nuclear symmetry energy as a function of density. Lower window: the corresponding proton fractions in neutron stars at β equilibrium.



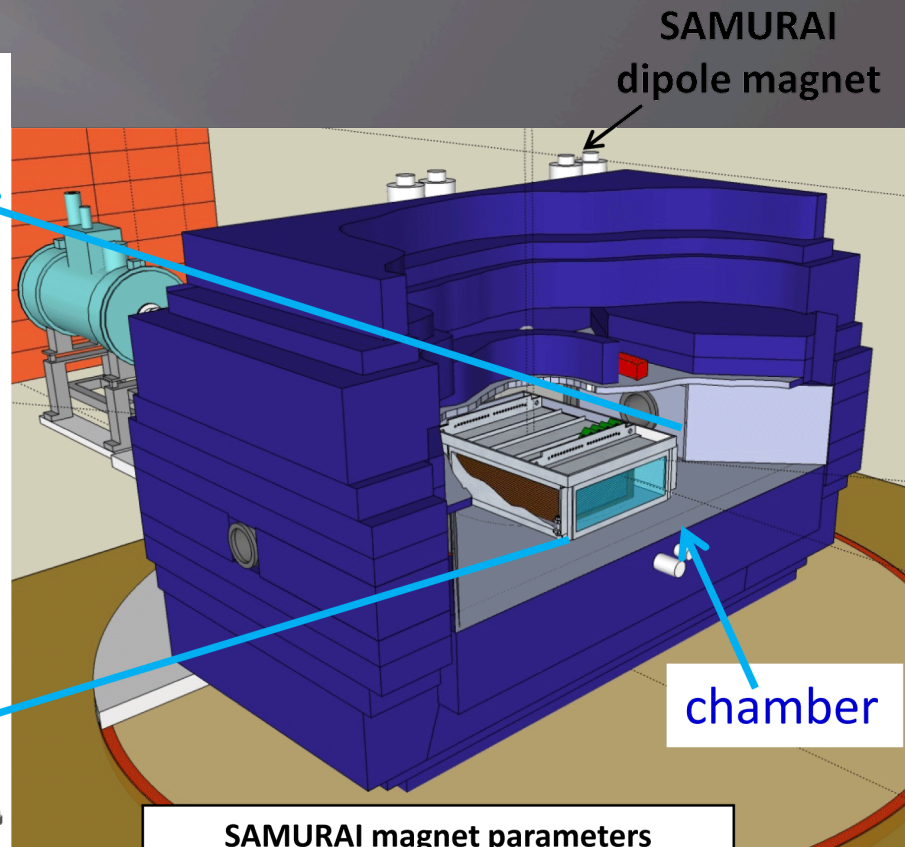
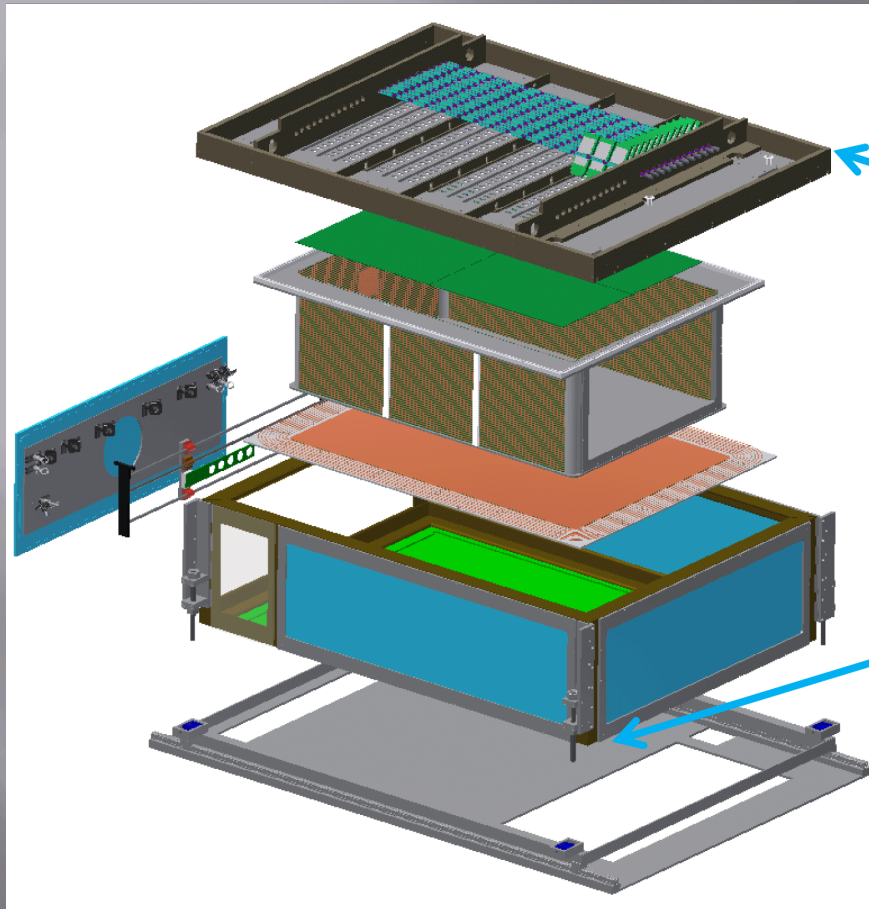
$$\pi^- / \pi^+ \equiv (5N^2 + NZ) / (5Z^2 + NZ)$$

Influence of symmetry pressure; Possible Probe

The symmetry pressure expels neutrons from and attracts protons to high density region of neutron-rich system. This suppresses $Y(n)/Y(p)$, $Y(\pi^-)/Y(\pi^+)$, etc.

- ▣ $\pi^+-\pi^-$ ratio
- ▣ Proton-neutron ratio
- ▣ Light ion ratio (t - ^3He)
- ▣ Particle flow of pions, protons, neutrons and light ions

Construction of SPiRIT



SAMURAI magnet parameters

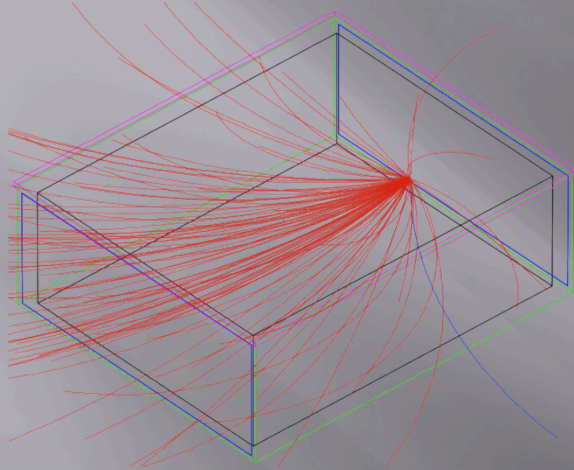
$B_{\text{typ}}, B_{\text{max}}$	0.5T, 3T
R, pole face	1 m
Gap	80 cm
Usable gap	75 cm

Supported by USA DoE funding (\$1.2M), and Japanese Grant-in-Aid for Scientific Research on innovative areas (\$1.3M).

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Star Matter

2013/10/25

Characteristics of SPiRIT(TPC?)



GEANT $^{132}\text{Sn}+^{124}\text{Sn}$ simulation collisions at $E/A=300$ MeV

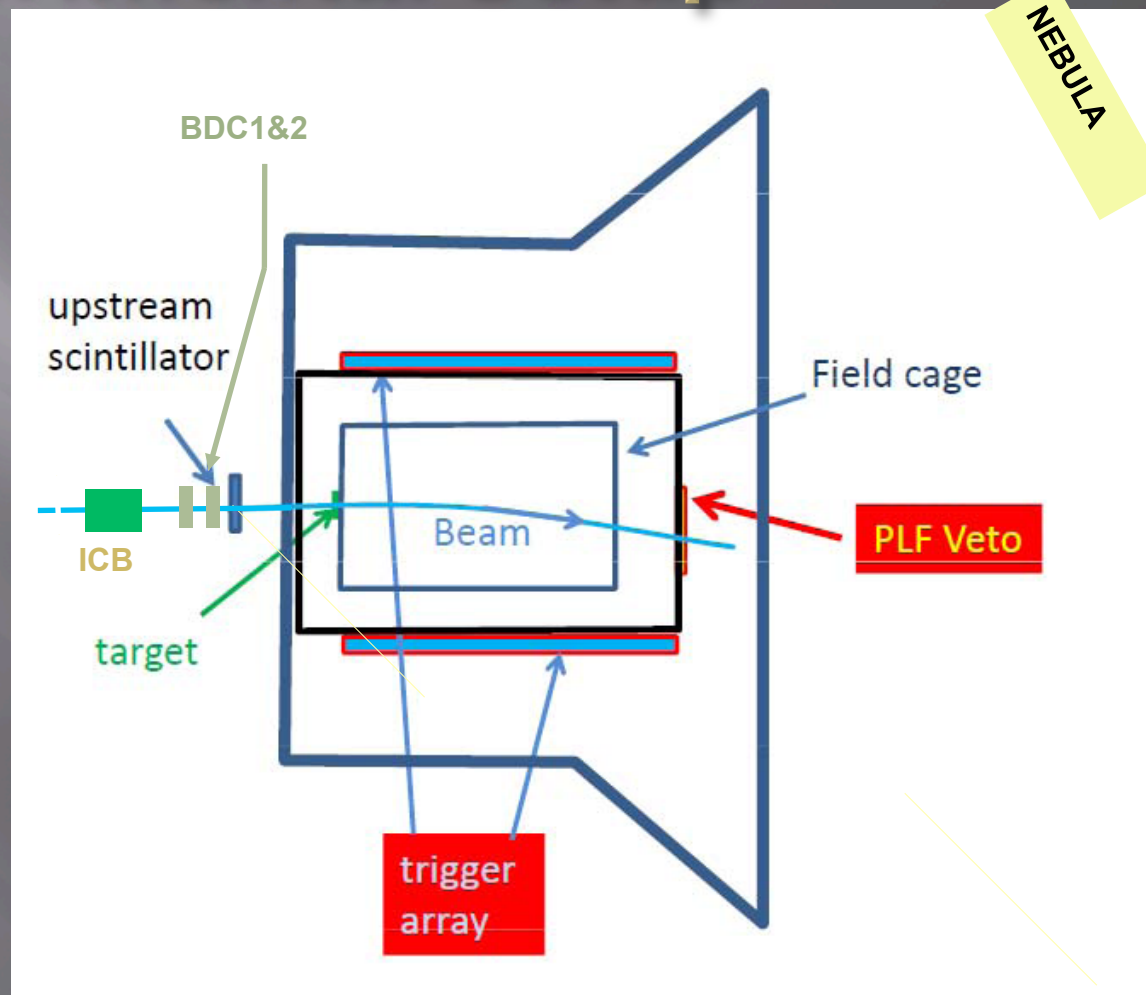
- Good efficiency for pion track reconstruction is essential.
- Initial design is based upon EOS TPC, whose properties are well documented.

SAMURAI TPC parameters	
Pad plane area	1.34m x 086 m
Number of pads	12096 (108 x 112)
Pad size	12 mm x 8 mm
Drift distance	53 cm
Pressure	1 atmosphere
dE/dx range	Z=1-3 (STAR El.), 1-8 (GET El.)
Two track resolution	2.5 cm
Multiplicity limit	200 (may impact absolute pion eff. in large systems.)

Experimental Setup

Equipment

- TPC - measures:
 - π^+ , π^- , p, d, t, ^3He , ^4He , IMF's
 - The SAMURAI chamber is at air
- Trigger scint. array:
 - selects central collisions and suppresses peripheral collisions.
- NEBULA:
 - provides neutron information

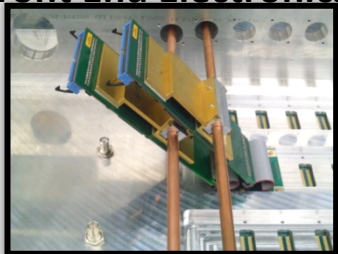


Magnet is at 0 degrees.

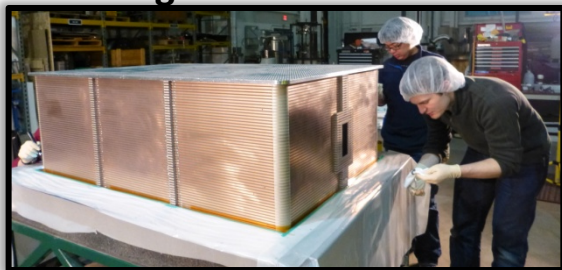
SPiRiT: Exploded View

Samurai **P**ion-**R**econstruction and **I**on-tracker **T**PC

Front End Electronics



Field Cage

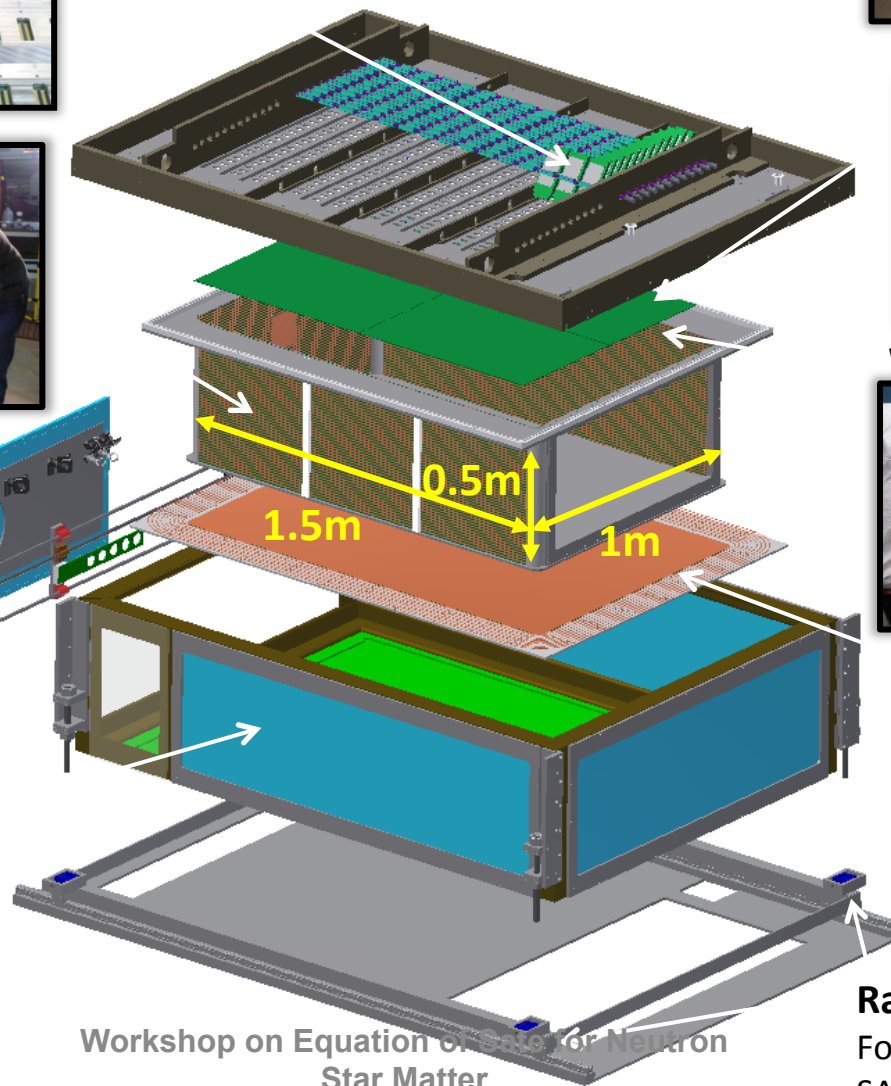


Beam

Calibration
Laser Optics

Target Mechanism

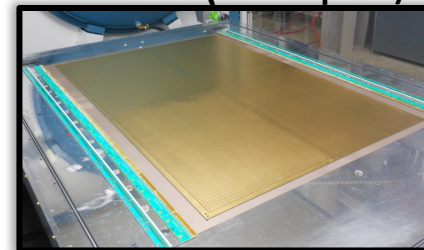
Thin-Walled Enclosure



Rigid Top Plate



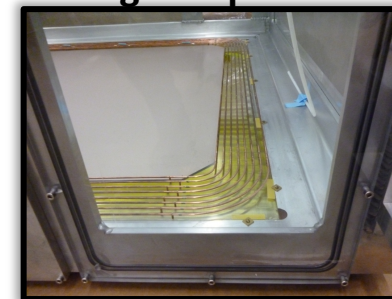
Pad Plane (12096 pads)



Wire Planes (e- mult)



Voltage Step-Down



Rails

For inserting TPC into
SAMURAI vacuum chamber

Read-out Electronics; GET

State of Arts
technology

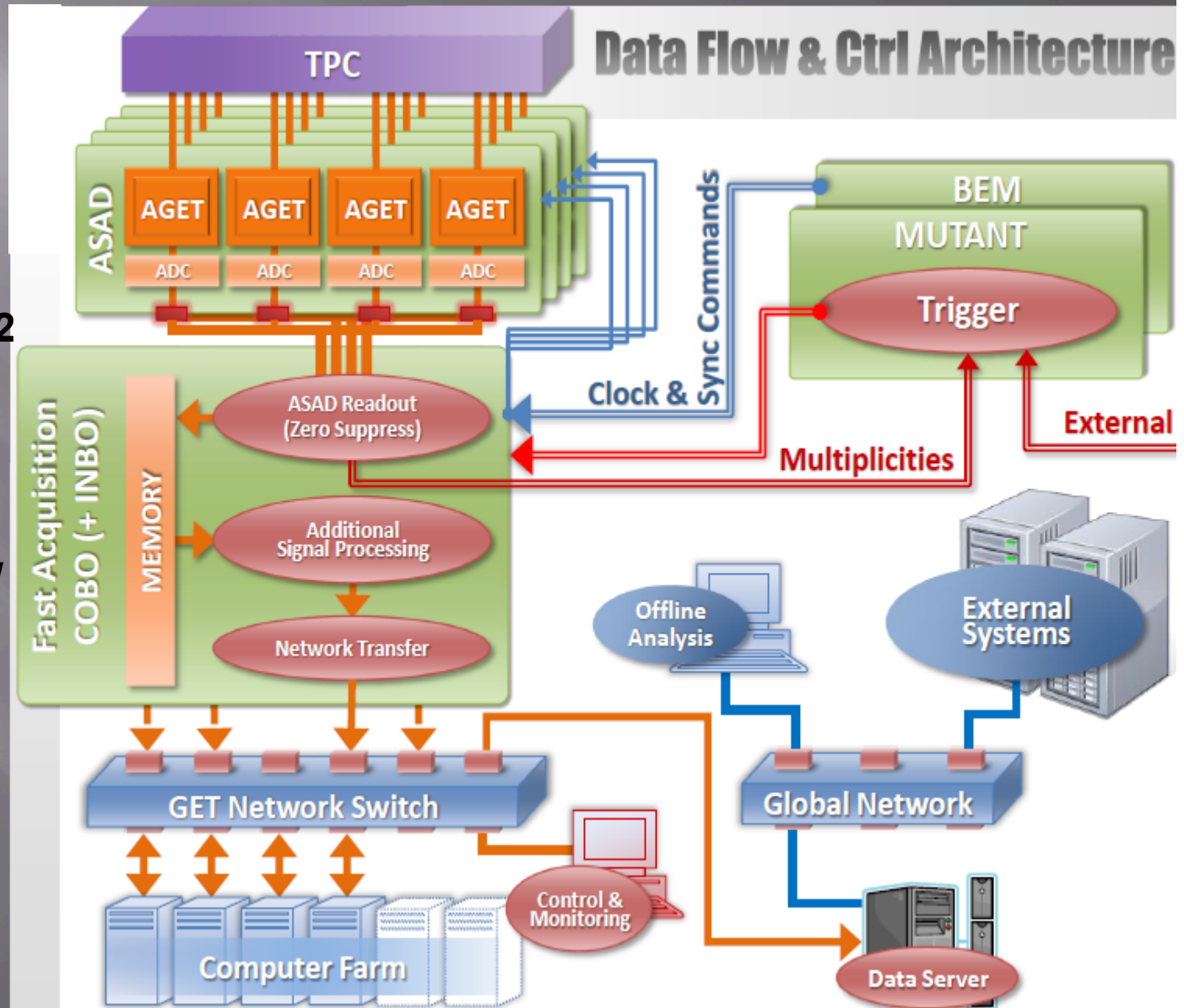
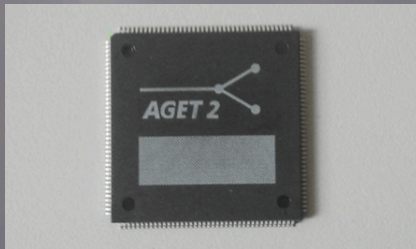
Capable to handle
1KHz – 10Gb/s

Wide dynamic range
10.5 bits

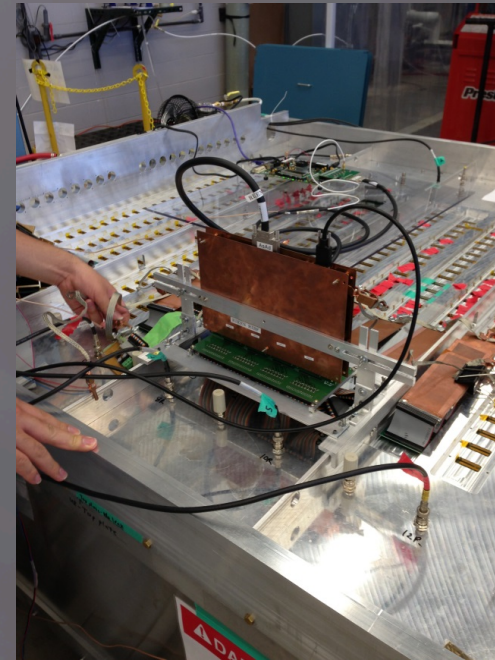
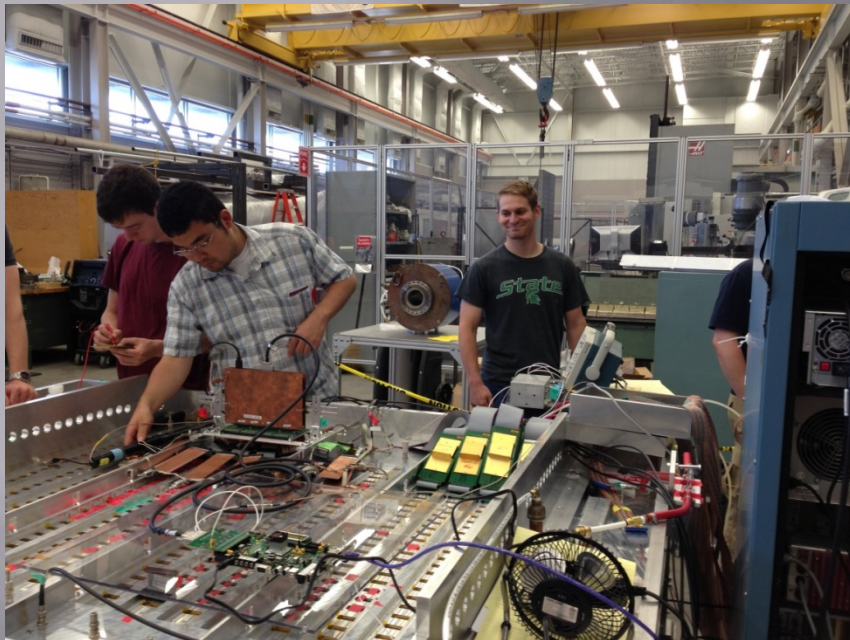
Capacitive Array 1-512

Sampling 1-100MHz

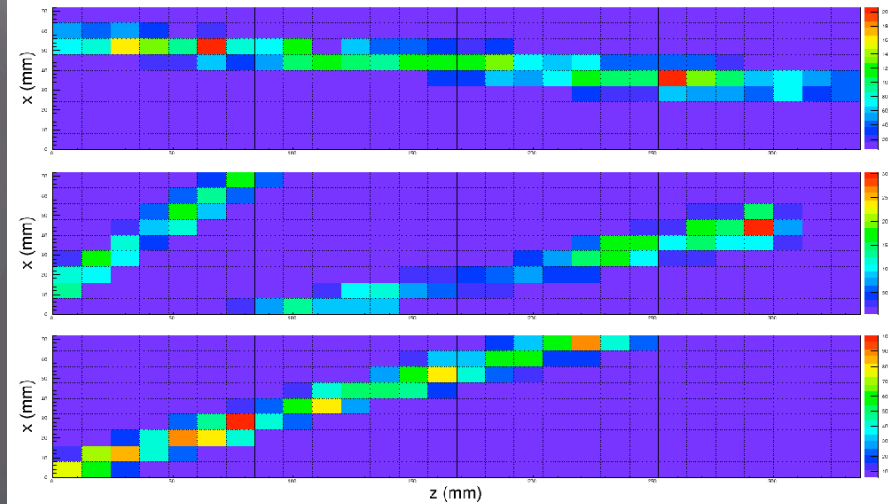
CEA/Bordeaux/GANIL/
MSU Collaboration



GET on SPiRIT



Event Displays on Pad Plane

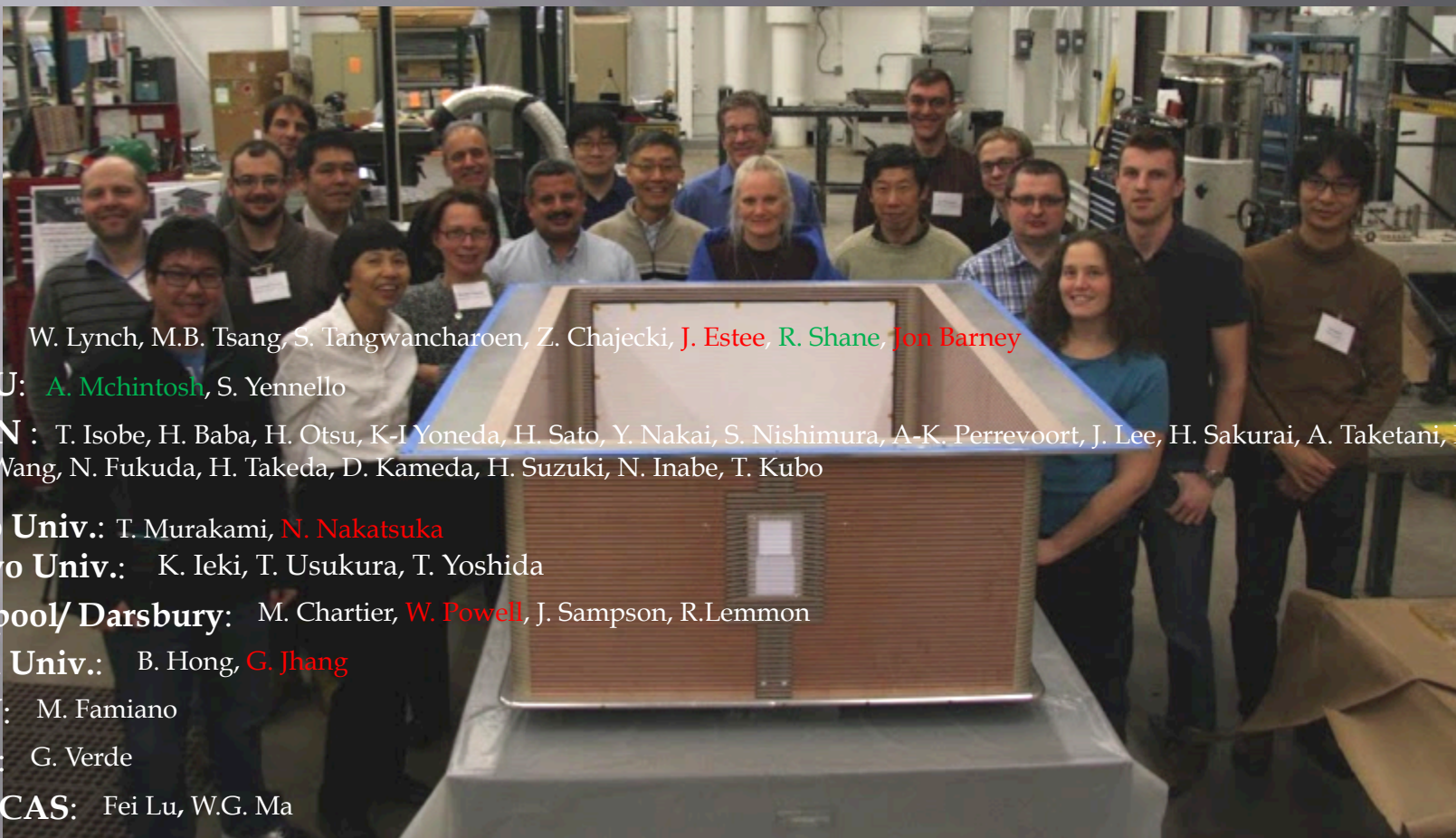


CoBo_2013-07-27T15_35_00.669_0000.graw

Summary

- Consistent constraints on the symmetry energy at sub-saturation densities with different experiments have been obtained.>> Heavy ion collisions should provide a good probe at high density..
- Astronomical observations suggest the importance of probing $\sim 2\rho_0$ region.
- We will be ready to measure pion and light fragment flow from $^{132}\text{Sn}+^{124}\text{Sn}$ and $^{124}\text{Sn}+^{112}\text{Sn}$ collisions at 300 MeV/u using the TPCs at RIKEN RIBF to constraint the symmetry energy above saturation densities by fall of 2014.

SPiRiT Collaboration



MSU: W. Lynch, M.B. Tsang, S. Tangwanchaoen, Z. Chajecski, **J. Estee**, **R. Shane**, **Jon Barney**

TAMU: **A. Mchintosh**, S. Yennello

RIKEN : T. Isobe, H. Baba, H. Otsu, K-I Yoneda, H. Sato, Y. Nakai, S. Nishimura, A-K. Perrevoort, J. Lee, H. Sakurai, A. Taketani, He Wang, N. Fukuda, H. Takeda, D. Kameda, H. Suzuki, N. Inabe, T. Kubo

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Rikkyo Univ.: K. Ieki, T. Usukura, T. Yoshida

Liverpool/Darbury: M. Chartier, **W. Powell**, J. Sampson, R. Lemmon

Korea Univ.: B. Hong, **G. Jhang**

WMU: M. Famiano

INFN: G. Verde

SIAP CAS: Fei Lu, W.G. Ma

CEA: E. Pollacco

INP: J. Lukasik

ORNL: A. Galindo-Uribarri

Tohoku Univ.: T. Kobayashi

TITech: T. Nakamura, Y. Kondo



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